

A DEEP SUB-MILLIMETER SURVEY OF LENSING CLUSTERS: A NEW WINDOW ON GALAXY FORMATION AND EVOLUTION

IAN SMAIL^{1,2} R. J. IIVISON^{2,3} & A. W. BLAIN⁴

¹) Department of Physics, University of Durham, South Rd, Durham DH1 3LE, UK

³) Institute for Astronomy, Dept. of Physics & Astronomy, University of Edinburgh, Blackford Hill, Edinburgh EH9 3HJ, UK

⁴) Cavendish Laboratory, Madingley Rd, Cambridge CB3 0HE, UK

ACCEPTED VERSION

ABSTRACT

We present the first results of a sub-millimeter survey of distant clusters using the new Sub-mm Common-User Bolometer Array (SCUBA) on the James Clerk Maxwell Telescope. We have mapped fields in two massive, concentrated clusters, A370 at $z = 0.37$ and Cl2244-02 at $z = 0.33$, at wavelengths of 450 and 850 μm . The resulting continuum maps cover a total area of about 10 arcmin 2 to 1σ noise levels less than 14 and 2 mJy beam $^{-1}$ at the two wavelengths, 2-3 orders of magnitude deeper than was previously possible. We have concentrated on lensing clusters to exploit the amplification of *all* background sources by the cluster, improving the sensitivity by a factor of 1.3-2 as compared with a blank-field survey. A cumulative source surface density of $(2.4 \pm 1.0) \times 10^3$ degree $^{-2}$ is found to a 50% completeness limit of ~ 4 mJy at 850 μm . The sub-mm spectral properties of these sources indicate that the majority lie at high redshift, $z > 1$. Without correcting for lens amplification, our observations limit the blank-field counts at this depth. The surface density is 3 orders of magnitude greater than the expectation of a non-evolving model using the local *IRAS* 60 μm luminosity function. The observed source counts thus require a substantial increase in the number density of strongly star-forming galaxies in the high-redshift Universe and suggest that optical surveys may have substantially underestimated the star formation density in the distant Universe. Deeper sub-mm surveys with SCUBA should detect large numbers of star-forming galaxies at high redshift, and so provide strong constraints on the formation of normal galaxies.

Subject headings: cosmology: observations — cosmology: early universe — galaxies: evolution — galaxies: formation — gravitational lensing — radio continuum: galaxies

1. INTRODUCTION

Surveys of the local Universe have shown that a third of the total bolometric luminosity is emitted at sub-mm and far-infrared wavelengths as a result of reprocessing of star-light by dust (Soifer & Neugebauer 1991). Moreover, some of the most vigorous star-forming galaxies in the local Universe are also those in which the effects of dust obscuration are most significant. While there have been striking advances in the identification of ‘normal’ galaxies at high redshift ($z \sim 2$ -4) using Lyman-dropout techniques (Steidel et al. 1996), such approaches are insensitive to highly obscured star-forming galaxies at these epochs. The presence of at least modest amounts of dust in distant protogalaxies, especially forming spheroids, is expected given the highly metal-enriched ISM which must be present during their formation (e.g. Mazzei & de Zotti 1996). Thus techniques sensitive to obscured distant galaxies should be used to investigate the formation of these populations in the early Universe.

Sensitive sub-mm observations present the first opportunity to detect dusty star-forming galaxies at high redshift. At wavelengths around 100 μm , the bulk of the luminosity of normal, star-forming galaxies is reprocessed star-light from dust and so observations in the sub-mm band can provide robust estimates of both the dust mass and total star formation rate in a galaxy. Furthermore, the negative K -correction at wavelengths $\lambda \gtrsim 400 \mu\text{m}$ means that sub-mm observations select star-forming galaxies at $z \gtrsim 1$ in an almost distance-independent manner, providing an efficient method for the detection of star-forming galaxies at very large redshifts, $z \lesssim 10$, and hence studying galaxy

evolution (Blain & Longair 1993, 1996 — BL96; Blain 1997; Eales & Edmunds 1997; Franceschini et al. 1997; Guideroni et al. 1997).

Most published sub-mm studies of distant galaxies have targeted atypical galaxies (e.g. radio-loud galaxies, Ivison et al. 1998). We report here the first deep sub-mm survey to probe the nature of normal galaxies at moderate and high redshift, $z \gtrsim 0.5$. In this study we have attempted to maximise the available sample of distant galaxies by concentrating on fields in moderate-redshift clusters. While the dominant spheroidal populations of these clusters are expected to be quiescent in the sub-mm band, the in-fall of field galaxies associated with the growth of the clusters (Smail et al. 1997) means that these fields may contain over-densities of moderate-redshift star-forming galaxies, as compared with ‘blank’ field surveys.

The main attraction of the clusters observed here, however, is that they are strong gravitational lenses, magnifying any source lying behind them (Blain 1997). Given the expected steep rise in the sub-mm counts (BL96), amplification bias could increase the source counts in these fields by a substantial factor, with a maximum predicted surface density of about 10 sources per SCUBA field down to 1 mJy at 850 μm (Blain 1997). Moreover, by targeting those clusters that contain giant arcs, images of distant field galaxies magnified by factors of 10-20, we can also obtain otherwise unachievable sensitivity ($\lesssim 0.1$ mJy at 850 μm) on the dust properties of a few serendipitously-positioned normal galaxies at high redshift. The angular scales of the region where these highly magnified high-redshift galaxies are found is also well-matched to the SCUBA field-of-view.

²PPARC Advanced Fellow.

Table 1
SCUBA observations of A370 and Cl 2244–02

Target	R.A. (J2000)	Dec. (J2000)	λ (μm)	Exposure time (ks)	Area (arcmin 2)	Flux density 1σ (mJy)
Cl 2244–02	22 47 11.9	–02 05 38	450	23.0	4.00	14.0
			850	23.0	5.40	1.9
A370	02 39 53.0	–01 35 06	450	25.7	4.35	13.3
			850	25.7	5.40	1.8

In the following sections we give details of the observations and their reduction, and discuss the results within the framework of current theoretical models of galaxy formation and evolution. We adopt $H_0 = 50 \text{ km s}^{-1} \text{ Mpc}^{-1}$ and $q_0 = 0.5$.

2. OBSERVATIONS AND REDUCTION

These data were obtained using SCUBA (Cunningham & Gear 1994) on the James Clerk Maxwell Telescope (JCMT)⁵. SCUBA contains a number of detectors and detector arrays cooled to 0.1 K and covering the atmospheric windows from $350 \mu\text{m}$ to $2000 \mu\text{m}$. In our survey, we operated the 91 element Short-wave (SW) array at $450 \mu\text{m}$ and the 37 element Long-wave (LW) array at $850 \mu\text{m}$, giving half-power beam widths of 7.5 and 14.7 arcsec respectively. Both arrays have a 2.3 arcmin instantaneous field-of-view and the design of the optics ensures that, with a suitable jiggle pattern for the secondary mirror,, fully sampled maps can be obtained simultaneously at 450 and $850 \mu\text{m}$. The multiplexing and high efficiency of the arrays means that SCUBA offers a gain in mapping speed of a factor of about 300 as compared with previous detectors.

The observations employed a 64-point jiggle pattern, fully sampling both arrays over a period of 128 s. The pattern was subdivided so that the target position could be switched between the signal and reference every 32 s: a repeating signal–reference–reference–signal scheme, with sixteen 1 s jiggles in each. Whilst jiggling, the secondary was chopped at 6.944 Hz by 60 arcsec in azimuth. The pointing stability was checked every hour and regular sky-dips were performed to measure the atmospheric opacity. The rms pointing errors were below 2 arcsec, while atmospheric zenith opacities at 450 and $850 \mu\text{m}$ were very stable during the course of each night, the night to night variations being in the range 0.98–2.2 and 0.18–0.37 respectively.

The dedicated SCUBA data reduction software (SURF, Jenness 1997) was used to reduce the observations. The reduction consisted of subtracting the reference from the signal after carefully rejecting spikes and data from noisy bolometers. Six quiet bolometers at the edge of each array were used to compensate for spatially-correlated sky emission. This reduced the effective noise-equivalent flux density from 100–350 to $90 \text{ mJy Hz}^{-1/2}$ at $850 \mu\text{m}$ (c.f. Ivison et al. 1998). The resulting maps were flatfielded, corrected for atmospheric attenuation, and calibrated using nightly beam maps of Uranus. The calibrated maps from each night were coadded and then linearly interpolated onto an astrometric grid using an approximately Nyquist sampling, of 2 and 4 arcsec pixel $^{-1}$ at 450 and $850 \mu\text{m}$ respectively, to produce the maps presented in Fig. 1 (Plate 1).

The final on-source integration times are listed in Table 1, along with field positions and sensitivity limits. The beams have moderate error lobes, especially at $450 \mu\text{m}$,

and so we correct our aperture measurements for flux outside the aperture using the measured values from our calibration sources. Even without including a factor to account for the lensing amplification, the data shown in Fig. 1 are the deepest sub-mm maps ever published, and illustrate the cosmetically clean and flat maps achievable with SCUBA in long integrations.

3. ANALYSIS AND RESULTS

Source catalogs from our fields were constructed using the SExtractor package (Bertin & Arnouts 1996). The detection algorithm requires that the surface brightness in 4 contiguous pixels exceeds a threshold (chosen as $\sim 1\sigma$ of the sky noise, Table 2), after subtracting a smooth background signal and convolving the map with a 4×4 pixel top-hat filter. The numbers of objects (N) detected in each field are given in Table 2, indicating that our observations are far from being limited by confusion: there are ~ 60 beams per source.

To assess the contribution of noise to our catalogs we re-ran the detection algorithm on the negative fluctuations in the map. This gave a simple estimate of the number of false-positive detections that may arise from the noise, assuming that the noise properties of the map are Gaussian. We estimate that there are no false detections in our catalogs (N_{-ve} , Table 2), and so all the detections are real. The presence of the brightest source in the reference beams (60 arcsec to the East and West in the $850 \mu\text{m}$ map of A370) was disregarded. This detection does confirm, however, the reality of positive features at this faint level, while the absence of any other negative detections limits the number of luminous sources which can lie in the regions covered by the reference beams.

Secondly, to determine the completeness of our sample we added a faint source to the maps repeatedly, re-ran our detection algorithm and estimated the efficiency of detecting this source as a function of its flux density. This provides a reliable estimate of the visibility of a faint compact source in the maps. The template source was a scaled version of our calibration source, Uranus. The estimated 80% and 50% completeness limits of the catalogs derived from these simulations are listed in Table 2 as $S_{80\%}$ and $S_{50\%}$. The incompleteness limits are relatively bright for the $450 \mu\text{m}$ maps because a large proportion of the flux density (about 40%) is found in the low surface brightness wings of the beam. The simulations also indicate that the measured $850 \mu\text{m}$ flux densities are unbiased and are typically accurate to 10% at 25 mJy and 30% at 4 mJy.

We discuss the detailed properties of the sources in another paper, where we also place limits on the dust masses of the numerous strongly-lensed distant galaxies covered by our maps. However, we note that, based on their weak or non-detection at $450 \mu\text{m}$, all of the $850 \mu\text{m}$ sources in

⁵The JCMT is operated by the Observatories on behalf of the UK Particle Physics and Astronomy Research Council, the Netherlands Organization for Scientific Research and the Canadian National Research Council.

TABLE 2
SUB-MM SOURCE COUNTS

Target	λ (μm)	Threshold (mJy beam^{-1})	Min. Area (arcsec^2)	N	N_{-ve}	$S_{80\%}$ (mJy)	$S_{50\%}$ (mJy)
Cl 2244-02	450	16.0	16	0	0	135	63
	850	1.9	64	2	0	5.5	3.9
A370	450	16.0	16	1	0	140	84
	850	1.8	64	4	0	5.4	3.9

our sample appear to have the sub-mm spectral characteristics of distant ($z \gtrsim 1$) star-forming galaxies and are thus unlikely to be associated with the clusters.

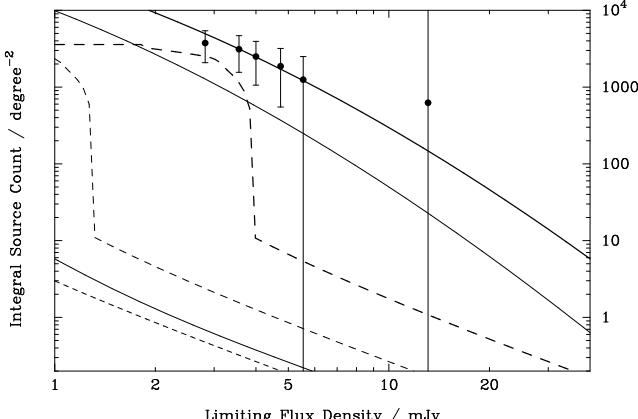


Figure 2 Models of the integral number counts of sources at $850\mu\text{m}$ in a parametric model of galaxy evolution (BL96) and from a simple model based on the limits on strongly star-forming systems in optical surveys of distant galaxies. The observations are represented by filled circles with Poisson errors, note that the errors are not independent on the various points. The observations have been corrected for the effects of lens amplification using simple lens models, but not corrected for incompleteness. The solid curves represent, in order of increasing predicted counts, models that include: no evolution; $(1+z)^3$ evolution with $z_{\text{max}} = 2$ and $z_0 = 5$ (Model A); and $(1+z)^3$ evolution with $z_{\text{max}} = 2.6$ and $z_0 = 7$ (Model B). The dashed lines represent models where we fill the Universe across $z = 0\text{--}10$ with a constant density ($0.6 \times 10^{-4} \text{ Mpc}^{-3}$) of star-forming galaxies with fixed star formation rates. In order of increasing predicted counts the dashed lines represent star formation rates for the population of: $M = 20, 50$ and $150 \text{ M}_\odot \text{ yr}^{-1}$, where we have assumed a dust temperature of $T = 60 \text{ K}$. Clearly only models including high densities of strongly star-forming galaxies are compatible with the observed surface density of sources.

Converting the observed number of sources into a surface density and correcting for incompleteness, we determine a cumulative number density across our two fields of $(2.4 \pm 1.0) \times 10^3 \text{ degree}^{-2}$ down to a 50% completeness limit of 4 mJy at $850\mu\text{m}$ (all errors include only Poisson contributions). At $450\mu\text{m}$, the single source we detect places only a very weak limit on the likely surface density of the order of $1 \times 10^3 \text{ degree}^{-2}$ brighter than 80 mJy .

We now estimate the likely lens amplification factors, and so place tighter limits on the typical blank-field counts. Because the distances to the detected sources are unknown, this estimate will, by necessity, be crude and so we have not attempted a detailed analysis. The cluster potentials are modelled as isothermal spheres with masses and centers determined from the redshifts and observed shapes of the giant arc in each cluster (Kneib et al. 1993; Smail et al. 1996). In these models the mean amplification factors for typical background sources ($z \gtrsim 1$) are about 2 and 1.3 in the regions covered by our maps of A370 and Cl2244-02 respectively, while the observed area of the maps (5.4 arcmin^2 at $850\mu\text{m}$) corresponds to 1.8 and 4.0 arcmin^2 in the respective source planes. Correcting the flux densities of our sources to take account of the probable lens amplifications, but not correcting for incompleteness, we predict the source counts presented in Fig. 2.

These indicate integrated number densities in blank fields of $(2.5 \pm 1.4) \times 10^3$ (similar to the observed value due to the source counts having a form close to $S^{0.4}$) and $(3.6 \pm 1.6) \times 10^3 \text{ degree}^{-2}$ to flux limits of 4 and 3 mJy respectively at $850\mu\text{m}$.

From the flux densities associated with the resolved sources in the fields we calculate lower limits to the background radiation intensities of 2.6 and $2.4 \times 10^{-10} \text{ W m}^{-2} \text{ sr}^{-1}$ at wavelengths of $450\mu\text{m}$ and $850\mu\text{m}$ respectively, averaged over both fields. By extrapolating the $850\mu\text{m}$ counts using our best-fit model (Fig. 2 and §4) to faint flux densities, we estimate total intensities of extragalactic background radiation from discrete sources of about $26\text{--}28$ and $4.4\text{--}6.7 \times 10^{-10} \text{ W m}^{-2} \text{ sr}^{-1}$ at wavelengths of 450 and $850\mu\text{m}$ respectively. These background radiation intensities are broadly consistent with the tentative detection of an isotropic component of the background radiation in the sub-mm by Puget et al. (1996), who inferred $\nu I_\nu = (23 \pm 20)$ and $(2.7 \pm 2.0) \times 10^{-10} \text{ W m}^{-2} \text{ sr}^{-1}$ at wavelengths of 450 and $850\mu\text{m}$ respectively. If we assume that all of the background radiation intensity we infer is due to the formation of massive stars, then we expect that a density parameter of heavy elements of $\lesssim 6 \times 10^{-4}$ will have accumulated in the Universe by the present epoch. This density corresponds to about 1.1% of the density parameter in baryons if $\Omega_b = 0.05$, and so it is fully consistent with present limits. We reiterate, however, that these are tentative estimates, the accuracy of which depends on the models assumed for both the lens and the form of the counts of distant galaxies.

4. DISCUSSION

Due to the negative K -corrections expected for distant galaxies, sub-mm observations provide a good estimate of the volume density of luminous star-forming galaxies at $z \gtrsim 1$. In the absence of redshifts for all the sources, the evolution of this population can be understood by comparing parameterised models to the counts (BL96). The BL96 models are based on the $60\mu\text{m}$ luminosity function of *IRAS* galaxies (Saunders et al. 1990) and assume that the luminosities evolve as $(1+z)^3$ out to a redshift, z_{max} , and then maintain the enhanced luminosity out to a cutoff redshift, z_0 . The form of this evolution is motivated by the observations of similar behaviour in both the radio galaxy and QSO number counts (Dunlop & Peacock 1991) as well as the luminosity density of field galaxies at $z < 1$ (Lilly et al. 1996). BL96 also give predicted counts for a non-evolving model using the same luminosity function. The adopted parameters for the models described in that paper give source counts which roughly span the range predicted by other similar works (e.g. Guiderdoni et al. 1997). In Fig. 2, we plot both the no-evolution case and two other parametric models based on BL96: model A – Model 2 in BL96 – with values of $z_{\text{max}} = 2$ and $z_0 = 5$; and model B, has $z_{\text{max}} \simeq 2.6$ and $z_0 = 7$, although most combinations of $z_{\text{max}} \simeq 2.2\text{--}2.9$ and $z_0 \gtrsim 5$ give comparable results. Model

B was used to estimate the total background radiation intensity discussed in §3.

From Fig. 2 it can be seen that the no evolution predictions fall short by 2–3 orders of magnitude of the observations. Thus, this first analysis of a deep sub-mm survey indicates that the number density of strongly star-forming galaxies, and so the mean star formation rate in the distant Universe is considerably larger than that seen locally. To estimate the extent of this evolution we assume that all the detected sources lie beyond the clusters. We then require strong evolution, of the form given in model B, out to $z > 2$ to fit the $850\mu\text{m}$ counts. For consistency, we check the predictions from model B for the observed counts at $450\mu\text{m}$; 0.7 sources are expected in the two fields, in agreement with the single detection.

We conclude from the $850\mu\text{m}$ counts that the integrated star formation rate in the Universe, as traced by the number density of the most luminous sources, must continue to rise out to beyond $z > 1$ ($z > 2$ in our parametric models), extending the trend observed at $z < 1$ (Lilly et al. 1996). For sources at $z \gtrsim 1$, the typical luminosity we infer is $L_{\text{FIR}} \sim 0.5\text{--}1.0 \times 10^{13}\text{L}_{\odot}$, with a star formation rate of $\dot{M} \gtrsim 100\text{--}300\text{M}_{\odot}\text{yr}^{-1}$. Using the observed surface density of these objects and assuming a constant space density of sources between $z = 1\text{--}5$, we would predict a number density of strongly star-forming galaxies of: $N(\dot{M} \gtrsim 150\text{M}_{\odot}\text{yr}^{-1}) \sim 1.2 \times 10^{-4}\text{ Mpc}^{-3}$, at $z \gtrsim 1$.

Limits on the number density of strongly star-forming galaxies at $z \sim 2\text{--}3.5$ have recently been published by Madau et al. (1996 — M96) on the basis of Lyman-dropout surveys. Their limit is $N(\dot{M} > 20\text{M}_{\odot}\text{yr}^{-1}) < 0.6 \times 10^{-4}\text{ Mpc}^{-3}$. We plot in Fig. 2 three models using this number density of sources to uniformly populate the volume from $z = 0\text{--}10$, but allowing the corresponding star formation rate to vary: $\dot{M} = 20, 50$ and $150\text{M}_{\odot}\text{yr}^{-1}$. These values represent that originally given by M96, the M96 limit corrected for dust extinction as suggested by Pettini et al. (1997), and a mean star formation rate closer to that needed to fit our observations. A galaxy population consistent with the limits from Madau et al. predicts a source density 3 orders of magnitude lower than that observed (Fig. 2). Even if the modest dust extinction proposed by Pettini et al. is included, we still under-predict the observed surface densities (unless the dust in this population is very cold, $T = 40\text{K}$, and they have very large dust masses). To match the observed surface den-

sity of $850\mu\text{m}$ sources we must significantly increase the mean star formation, either by further increasing the star formation rate associated with the optically-selected samples (c.f. Meurer et al. 1997) or by introducing a population of strongly star-forming, but highly obscured, distant galaxies missed in these samples. Deeper sub-mm surveys (BL96; Pearson & Rowan-Robinson 1996) are necessary to differentiate between these possibilities and hence provide an unbiased view of star formation in the distant Universe.

5. CONCLUSIONS

- We have presented the first sub-mm survey of the distant Universe, deep enough that we should detect the evolving galaxies predicted by current theoretical models, while at the same time covering a sufficiently large area to be statistically reliable. We derive cumulative source counts of $(2.4 \pm 1.0) \times 10^3$ degree $^{-2}$ down to 4mJy at $850\mu\text{m}$.
- The surface density of faint sources in the sub-mm far exceeds a simple non-evolving model using the locally observed $60\mu\text{m}$ galaxy luminosity function. Thus our observations require a substantial increase in the number density of strongly star-forming galaxies at $z \gtrsim 1$.
- Comparison of our observations with the predictions of simple parametric models implies that the luminosity density of the brightest sub-mm sources continues to increase out to $z > 1$. Models based upon the claimed properties of star-forming galaxies from optically-selected samples of distant galaxies (Madau et al. 1996) significantly underestimate the observed surface density of sub-mm sources. We suggest that such samples may be missing a considerable proportion of the star formation in dust-obscured galaxies at high redshift. We conclude that question of the evolution of the star formation density in the distant Universe and hence the epoch of galaxy formation is still very much open.

ACKNOWLEDGEMENTS

We thank the SCUBA development and commissioning team for providing an efficient and user-friendly instrument, and Ian Robson for his support, enthusiasm and drive to see the best science done with the JCMT and SCUBA. We wish to thank Richard Bower, Richard Ellis, Carlos Frenk, Wayne Holland, Tim Jenness, Jean-Paul Kneib and Malcolm Longair for useful conversations and help. We also thank an anonymous referee for comments that clarified the content and presentation of this paper.

REFERENCES

Bertin, E., & Arnouts, S. 1996, A&A, 117, 393
 Blain, A.W. 1997, MNRAS, in press
 Blain, A.W., & Longair, M. 1993, MNRAS, 264, 509
 Blain, A.W., & Longair, M. 1996, MNRAS, 279, 847 (BL96)
 Cunningham, C.R., & Gear, W.K., Duncan, W.D., Hastings, P.R., & Holland, W.S., 1994, SPIE Proc. Vol. 2198, 638
 Dunlop, J.S., & Peacock, J.A. 1991, MNRAS, 247, 19
 Eales, S.A., & Edmunds, M.G. 1997, MNRAS, 286, 732
 Franceschini, A., Andreani, P., & Danese, L. 1997, MNRAS, submitted, (astro-ph/9706022)
 Guiderdoni, B., Hivon, E., & Bouchet, F.R. 1997, preprint, (astro-ph/9707134)
 Ivison, R.J., et al. 1998, ApJ, in press, (astro-ph/9709124)
 Jenness, T. 1997, Starlink User Note 216.1,
 (http://www.jach.hawaii.edu/jcmt_sw/scuba/surf/)
 Kneib, J.-P., Mellier, Y., Fort, B., & Mathez, G. 1993, A&A, 273, 367
 Lilly, S.J., Le Fevre, O., Hammer, F., & Crampton, D. 1996, ApJ, 460, L1
 Madau, P., Ferguson, H.C., Dickinson, M.E., Giavalisco, M., Steidel, C.S., & Fruchter, A., 1996, MNRAS, 283, 1388 (M96)
 Mazzei, P., & de Zotti, G. 1996, MNRAS, 279, 535
 Meurer, G.R., Heckman, T.M., Lehnert, M.D., Leitherer, C., & Lowenthal, J., 1997, AJ, 114, 55.
 Pearson, C., & Rowan-Robinson, M., 1996, MNRAS, 283, 174
 Pettini, M., Steidel, C.C., Dickinson, M.E., Kellogg, M., Giavalisco, M., & Adelberger, K.L., 1997, preprint, (astro-ph/9707200)
 Puget, J.-L., Abergel, A., Bernard, J.-P., Boulanger, F., Burton, W.B., Désert, F.-X., & Hartmann, D., 1996, A&A, 308, L5
 Saunders, W., Rowan-Robinson, M., Lawrence, A., Efstathiou, G., Kaiser, N., Ellis, R.S., & Frenk, C.S., 1990, MNRAS, 242, 318
 Soifer, B.T., & Neugebauer, G., 1991, AJ, 101, 354
 Smail, I., Dressler, A., Kneib, J.-P., Ellis, R.S., Couch, W.J., Sharples, R.M., & Oemler, A. 1996, ApJ, 469, 508
 Smail, I., Edge, A.C., Ellis, R.S., & Blandford, R.D. 1997, MNRAS, in press
 Steidel, C.C., Giavalisco, M., Pettini, M., Dickinson, M., & Adelberger, K. L. 1996, ApJ, 462, L17

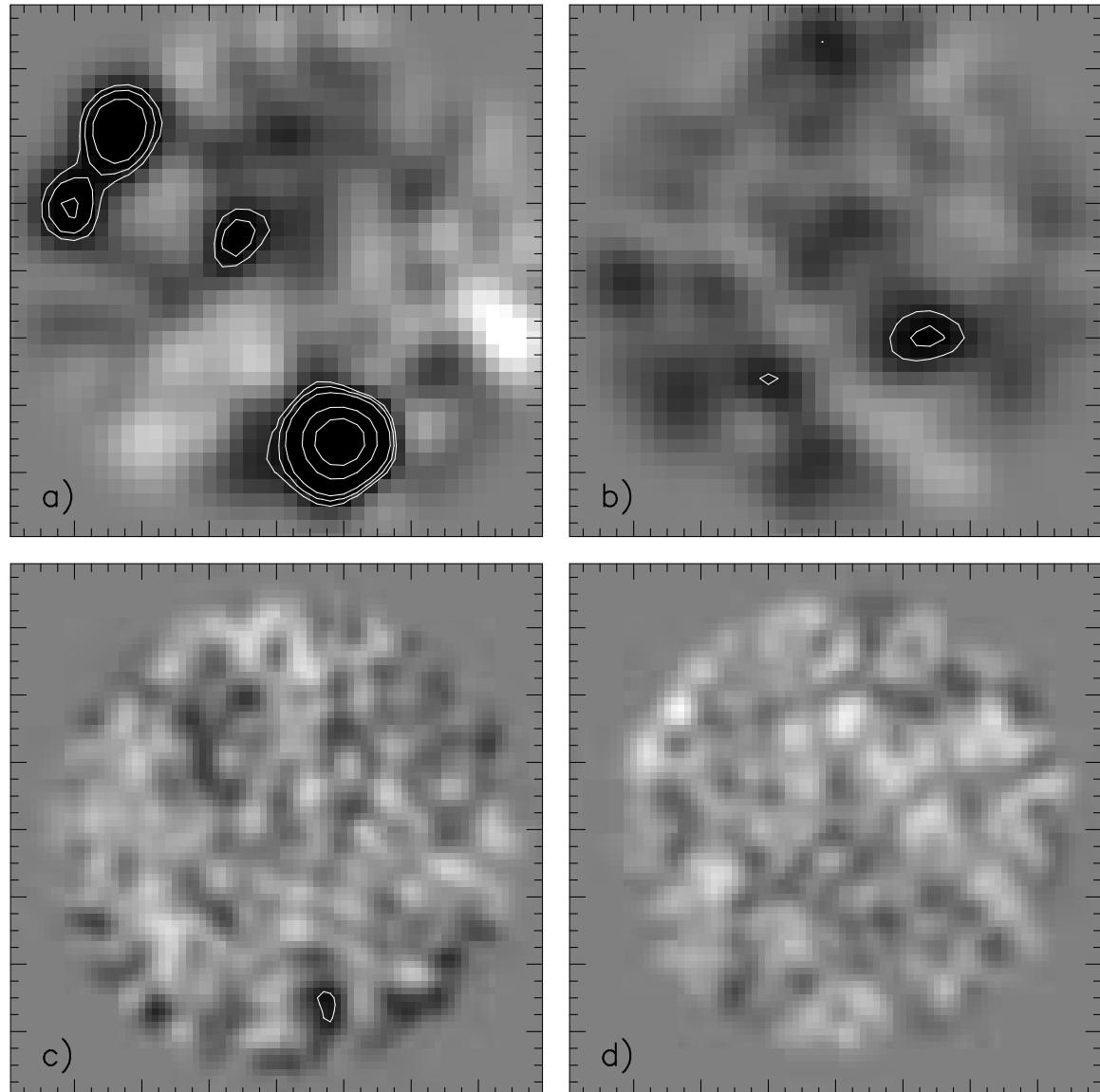


Figure 1 (Plate 1). The 450 and 850 μm maps of the two fields: a) A370, 850 μm ; b) Cl 2244-02, 850 μm ; c) A370, 450 μm ; d) Cl 2244-02, 450 μm . The maps are smoothed to the instrumental resolution at each wavelength and have had their boundaries apodized to remove regions on the outskirts with low effective exposure times resulting from the rotation of the field-of-view during our long exposures. They are displayed as a grayscale from -4σ to 4σ , the contours are positive and show 3, 4, 5, 10, 15 σ for each field. The major tick marks are 20 arcsec in all panels.